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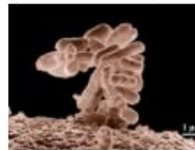
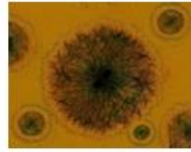
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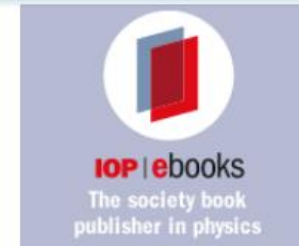
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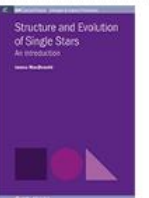
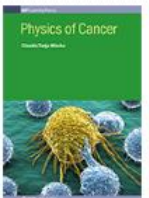


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In this book we will look at what planetary nebulae are, where they come from and where they go. We will discuss what mechanisms cause these beautiful markers of stellar demise as well as what causes them to form their variety of shapes. How we measure various aspects of planetary nebulae such as what they are made of will also be explored. Though we will give some aspects of planetary nebulae mathematical treatment, the main points should be accessible to people with only a limited background in mathematics. A short glossary of some of the more arcane astronomical terms is at the end of the book to help in understanding. Included at the end of each chapter is an extensive bibliography to the peer reviewed research on these objects and I would encourage the reader interested in an even deeper understanding to read these articles.

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**A**

**Author**

Jason J Nishiyama

**Published**

May 2018

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## An Introduction to Planetary Nebulae

CHAPTER 1 • FREE TO READ

### Introduction

Jason J Nishiyama

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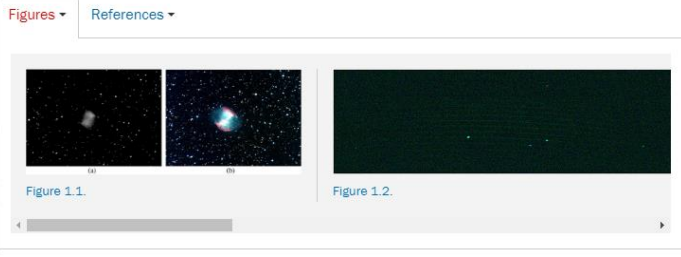
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### Abstract

In chapter 1 we look at what planetary nebulae are, when they were discovered and how they received their rather curious name. How we came to understand where planetary nebulae came from and how they evolve is also explored. How we look for these objects today is explained. Finally the components of these objects are described.



C



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Abstract

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1.3. Components of planetary nebulae

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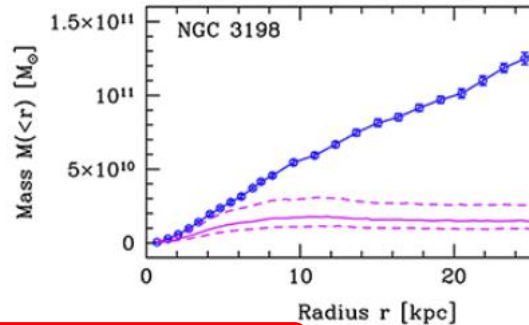
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Galaxy, it is the gravity of the system that ensures that they don't fly off into the Universe! The faster the stars rotate about the galactic centre, the more gravity, or mass, there must be to keep the stars bound, which can be seen by equating the centripetal force experienced by an orbiting star of mass  $m$  with the gravitational force;

$$\frac{mv^2}{r} = \frac{GM(<r)m}{r^2} \Rightarrow M(<r) = \frac{v^2 r}{G} \quad (1)$$

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**Figure 2.** (Left) Spiral Galaxy NGC 3198, imaged by the Sloan Digital Sky Survey. The distance from the centre to the edge of the visible disk of stars is roughly  $8 \text{ kpc} = 2.5 \times 10^{19} \text{ m}$ . (Right) The mass  $M(<r)$  enclosed within a radius  $r$  as inferred from the distribution of visible stars, i.e. the stellar mass (shown in pink with the uncertainty on this measurement indicated by the upper and lower dashed pink curves). This can be compared with the total mass,  $M(<r)$  inferred from the rotational velocities of stars (blue data points). As the radius increases, the total enclosed mass continues to grow. The stellar mass, in contrast, stops increasing at roughly  $8 \text{ kpc}$ , where the visible stellar disk ends. This discrepancy provides strong evidence to support the theory of dark matter. Data Source: Matthew Bershady and collaborators.

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